

The Eccentrids of the Mars family in the minor bodies system

Alexandra Terentjeva and Elena Bakanas

Institute of Astronomy of the Russian Academy of Sciences, Moscow, Russia

ater@inasan.ru; oterma@yandex.ru

A group of 39 Eccentridd asteroids ($e \geq 0.4$) belonging to the Mars family has been identified. Of the 670 orbits of meteoroid streams in several catalogues, only one Librid meteoroid stream, No. 50, has been found (Terentjeva, 1966), which also belongs to the Eccentridds of the Mars family. Among the 39 asteroids, 13 groupings have been discovered, comprising 33 asteroids, i.e. 84.6% of the asteroids in the system of the Eccentridds of the Mars family represent ‘organized’ matter. The distribution of asteroids by longitude of perihelion has been studied, and some features of this unique system of bodies are discussed.

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1 Introduction

The Eccentridds as a system of meteoroid bodies were discovered by A. K. Terentjeva (Galibina and Terentjeva, 1981), although she had recorded meteoroid streams of the Eccentridds much earlier, in 1967–1968, when identifying minor meteor streams (Terentjeva, 1967). In our work (Terentjeva and Barabanov, 2016), we define the Eccentridds as a special group of meteoroid bodies, the orbits of which fall within the limits of $a < 1$ AU, $e \geq 0.4$, $Q \leq 1.1$ AU. In this work, 52 asteroids were found among 11673 near-Earth objects (NEOs), and some features of this system of bodies were studied. Eight asteroid streams of the Eccentridds were found. Together with the Eccentridd meteoroid system, the Eccentridd asteroid system forms a single population of minor bodies of the Solar System.

But are there in fact Eccentridds of the Mars family, i.e. minor bodies with orbits of high eccentricity belonging to the Mars family? This is the question we set in this work.

2 Research results

The work of M. Bielicki (1972) provides the limit values of the perturbation region for the planets of the Solar System. The limiting distances for planetary families (AU):

- Mercury 0.002
- Venus 0.010
- Earth 0.014
- Mars 0.010
- Jupiter 0.50
- Saturn 0.62
- Uranus 0.67
- Neptune 1.11
- Pluto 0.55

These data show that the radius of the sphere of influence of Mars on the perturbed body is very small. Our selection criteria for the orbits of the Eccentridds of the Mars family were: $e \geq 0.4$ and $1.51 \leq Q \leq 1.53$ AU. From the NASA JPL catalogue¹, which contained 1448137 asteroids as of 11 May 2025, 39 asteroid orbits of the Eccentridds of the Mars family were selected (*Table 1*).

Table 1 – The Eccentridd asteroids of the Mars family. Equinox J2000.0. H is the absolute magnitude of the asteroid, D is the diameter of the asteroid.

No	Full name	Epoch	a (AU)	e	q (AU)	Q (AU)	Ω (°)	ω (°)	i (°)	π (°)	H	D km	Period years
2	2016 CD31	05/05/2025	1.044	0.454	0.570	1.519	240.49	12.71	0.46	253.20	26.90		1.067
3	2008 JC	05/05/2025	1.045	0.463	0.561	1.529	47.86	58.17	2.50	106.03	26.50		1.069
4	2013 SL20	05/05/2025	1.080	0.403	0.645	1.516	8.49	115.65	4.29	124.14	24.62		1.123
5	2013 AH76	01/17/2013	0.985	0.547	0.447	1.523	347.19	353.77	4.55	340.96	22.88		0.978
6	2014 JN2	05/03/2014	0.894	0.692	0.276	1.512	219.84	148.81	6.32	8.64	24.27		0.845
7	2002 JD9	05/05/2025	1.056	0.440	0.591	1.521	208.16	139.08	6.61	347.24	23.20		1.085
8	2018 VD	05/05/2025	1.068	0.419	0.620	1.516	38.78	252.38	7.61	291.17	27.10		1.103

¹ <https://ssd.jpl.nasa.gov/>

No	Full name	Epoch	a (AU)	e	q AU	Q AU	Ω (°)	ω (°)	i (°)	π (°)	H	D km	Period years
9	2022 DY1	05/05/2025	1.050	0.448	0.579	1.520	156.05	249.83	7.85	45.89	27.92	--	1.076
10	2018 BB5	05/05/2025	0.956	0.593	0.389	1.523	128.84	136.24	7.96	265.08	24.40		0.935
11	2018 KL2	05/05/2025	1.051	0.437	0.591	1.510	212.55	170.89	8.66	23.44	21.70		1.077
12	2019 EO	05/05/2025	0.994	0.532	0.465	1.523	130.26	255.55	9.35	25.80	23.18		0.991
13	2010 LJ61	05/05/2025	1.046	0.460	0.565	1.526	70.83	48.35	9.80	119.18	20.95	0.192	1.069
14	2021 VP6	05/05/2025	1.030	0.483	0.533	1.527	16.55	181.06	10.28	197.61	21.57	--	1.045
15	4034 Vishnu 1986 PA	05/05/2025	1.059	0.444	0.589	1.530	157.87	296.69	11.17	94.56	18.48	0.420	1.090
16	2021 TJ57	05/05/2025	1.008	0.507	0.497	1.519	4.68	180.12	11.63	184.80	21.46	--	1.012
17	2007 US	05/05/2025	0.962	0.583	0.402	1.523	24.02	203.06	12.40	227.08	22.74		0.944
18	516155 2016 DP	05/05/2025	1.022	0.477	0.535	1.510	54.36	254.00	13.33	308.36	20.11		1.034
19	152742 1998 XE12	05/05/2025	0.878	0.739	0.229	1.527	280.02	353.16	13.44	273.18	19.10	0.413	0.823
20	512234 2015 VO66	05/05/2025	0.952	0.600	0.381	1.523	64.46	219.31	14.02	283.77	20.62		0.929
21	2016 GG216	05/05/2025	1.008	0.515	0.489	1.528	171.94	157.16	14.71	329.10	23.15		1.013
22	2019 SM2	09/21/2019	0.980	0.550	0.440	1.519	174.65	51.02	15.08	225.67	26.10	--	0.969
23	2020 KK5	05/05/2025	0.932	0.634	0.341	1.523	83.25	325.07	15.15	48.32	21.93	--	0.900
24	88213 2001 AF2	05/05/2025	0.954	0.595	0.386	1.522	114.25	194.99	17.81	309.24	19.48	0.912	0.932
25	2014 MS67	05/05/2025	1.021	0.496	0.514	1.527	253.71	217.37	18.80	111.07	21.20		1.031
26	481127 2005 UJ6	05/05/2025	1.068	0.416	0.624	1.513	42.36	131.41	19.00	173.77	21.50		1.104
27	2025 DU2	05/05/2025	1.043	0.452	0.572	1.514	345.23	291.59	19.07	276.82	24.25	--	1.065
28	2022 FV	05/05/2025	1.062	0.428	0.607	1.516	3.90	300.62	19.34	304.53	24.15	--	1.094
29	3753 Cruithne 1986 TO	05/05/2025	0.998	0.515	0.484	1.511	126.19	43.88	19.80	170.07	15.44	2.071	0.997
30	234341 2001 FZ57	05/05/2025	0.944	0.604	0.374	1.515	22.09	339.98	20.67	2.07	18.93	0.344	0.918
31	2012 VN6	05/05/2025	1.043	0.461	0.563	1.524	216.94	26.24	21.03	243.19	21.38		1.065
32	141531 2002 GB	05/05/2025	0.992	0.529	0.467	1.517	40.81	8.32	22.55	49.13	19.15	0.303	0.988
33	2008 LC2	06/03/2008	1.012	0.499	0.507	1.518	78.92	303.73	23.65	22.65	24.25		1.018
34	2011 FP29	04/01/2011	1.055	0.447	0.584	1.526	9.08	332.76	24.79	341.84	21.95		1.083
35	2016 VU2	05/05/2025	0.987	0.548	0.446	1.527	235.37	18.54	26.79	253.91	22.00		0.980
36	2020 UF3	05/05/2025	0.967	0.574	0.412	1.522	29.65	127.45	27.49	157.09	28.38	--	0.951
37	2015 KB57	05/05/2025	0.945	0.607	0.371	1.518	239.61	208.83	29.67	88.44	22.30		0.918
38	250458 2004 BO41	05/05/2025	1.019	0.493	0.517	1.521	337.74	254.46	35.55	232.19	18.00	0.630	1.028
39	2023 VT10	11/16/2023	0.903	0.694	0.277	1.529	232.52	326.71	39.27	199.23	22.21	--	0.858

Table 2 – Orbital elements and other parameters of the Librid meteoroid stream (J1950.0)¹. The source for both is No 50 in Terentjeva (1966).

Date (UT)	α_g (°)	δ_g (°)	v_∞ km/s	a AU	e	q AU	Q AU	ω (°)	Ω (°)	i (°)	π (°)
1954 IV 7.30	206	−11	18.8	1.04	0.48	0.55	1.5	294	17	0	311
1953 IV 11.40	212	−10	18.1	1.02	0.46	0.55	1.5	296	21	2	317

(¹) Over 50 years, the difference due to precession is less than 1°, which is not significant.

Now let us consider meteoroid streams. We reviewed the most significant known catalogues of orbits of meteoroid and fireball streams (Cook, Lindblad, Terentjeva, Halliday et al., Rudawska, Jenniskens), totalling 670 orbits of streams. It turned out that among them there is only one meteoroid stream of the Librids, No. 50 from the catalogue by A. Terentjeva (1966), which meets the selected criteria and is an Eccentrid of the Mars family. The elements of its orbit and other parameters are given in Table 2. Thus, the Eccentrids of the Mars family constitute 0.0027% of asteroids and 0.15% of meteoroid streams, i.e. two orders of magnitude more.

Let us analyze the data for 39 asteroids of the Eccentrids of the Mars family (Table 1). Of the 39 asteroids, 8 asteroids, i.e. 20%, enter at perihelion inside the orbit of Mercury. These are asteroids with absolute magnitudes of 19–24 and diameters ranging from 344 m to 912 m (according to three estimates). The remaining 31 asteroids out of 39 enter at perihelion inside the orbit of Venus. These data indicate the role of these planets in gravitational influence. Let us consider the distribution of 39 asteroids by longitude of perihelion, π (Figure 1). Two maxima are clearly visible in the intervals $0^\circ < \pi < 30^\circ$ and $270^\circ < \pi < 300^\circ$ (12.8% of asteroids in each). It is also worth noting a slightly smaller third maximum (10.3% of asteroids) in the interval $90^\circ < \pi < 120^\circ$, on either side of which there is a deficit of longitudes of perihelion in the intervals from $\pi = 60^\circ$ to $\pi = 90^\circ$ and from $\pi = 120^\circ$ to $\pi = 150^\circ$.

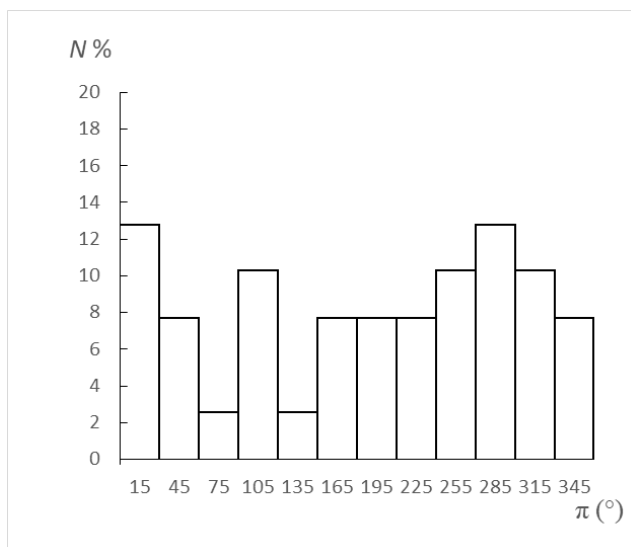


Figure 1 – Distribution of the longitudes of perihelion π of the Eccentrid asteroids of the Mars family.

I. I. Putilin (1953), studying the distribution of longitudes of perihelion of asteroids in the Main Belt, found a clear maximum in the interval $0^\circ < \pi < 20^\circ$. This is now well known and is explained by the action of secular perturbations from large planets, mainly Jupiter ($\pi = 15^\circ.4$). This maximum of the Main Belt lies within the slightly wider first maximum of the Eccentrids of the Mars family in the interval $0^\circ < \pi < 30^\circ$ (Figure 1). None of the nine known planets falls within the region of the second maximum of the Eccentrids of the Mars family in the interval from $\pi = 270^\circ$ to $\pi = 300^\circ$. However, directly

opposite the second maximum is the third maximum, with an interval from $\pi = 90^\circ$ to $\pi = 120^\circ$. Interestingly, Earth ($\pi = 102^\circ.8$) is located in this interval and appears to be involved in its formation.

As for the two deficits in longitudes of perihelion in the intervals $60^\circ < \pi < 90^\circ$ and $120^\circ < \pi < 150^\circ$ (Figure 1), Mercury falls into the first one ($\pi = 77^\circ.5$), and Venus into the second one ($\pi = 131^\circ.7$). Both planets have a very small radius of influence on the perturbed body, 0.002 AU for Mercury and 0.010 for Venus (see above).

I. I. Putilin discovered a dependence of the distribution of the longitude of perihelion on eccentricity. ‘The greater the eccentricity of the orbit, the more pronounced the tendency of perihelia to group around a certain longitude’, he wrote (Terentjeva and Barabanov, 2016). In the work mentioned, this dependence is confirmed on the system of the Eccentrid asteroids of the Earth group ($Q \leq 1.1$ AU). In our case with the Eccentrids of the Mars family, this dependence is even more pronounced (Figure 1). Among the population of 39 Eccentrid asteroids of the Mars family, orbits with $e = 0.5$ account for 44%, and those with $e = 0.6–0.7$ account for 26%. Apparently, on average, e of the Eccentrids of the Mars family is higher than that of the Eccentrids of the Earth group.

Similar to meteoroid bodies and asteroids of the system of the Eccentrids of the Earth group, the Eccentrid asteroids of the Mars family tend to make groupings. Among 39 asteroids (Table 1), we identified 15 groupings, of which 13 groupings have 2 orbits, one grouping has 3 orbits, and one has 4 orbits. These data are shown in Table 3.

Thus, 15 groupings include 33 orbits (out of 39). This means that 84.6% of the asteroid orbits in the system of the Eccentrids of the Mars family represent ‘organized’ matter.

Table 3 – Groupings of the Eccentrids of the Mars family.

Number, No. in Table 1	Number, No. in Table 1
1, 8	12, 33
2, 31	13, 25
3, 13, 15, 25	14, 16
4, 13	16, 29
7, 34	17, 22
1, 8, 28	18, 28
10, 35	31, 38
11, 34	

3 Conclusion

Mars, as we can see, is weak enough to create its own large family of minor bodies, since the radius of its gravitational influence is very small. Nevertheless, a group of 39 Eccentrid asteroids of the Mars family has been found, which together with the Eccentrids of the Earth group (Terentjeva and Barabanov, 2016) form a single exotic

population of minor bodies in the inner region of the Solar System.

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